

IMPACT OF PULSAR AND FALLBACK SOURCES ON MULTIFREQUENCY KILONOVA MODELS

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Draft version April 15, 2019

ABSTRACT

We explore the impact of pulsar electromagnetic dipole and fallback accretion emission on the luminosity of a suite of kilonova models. The pulsar models are varied over pulsar magnetic field strength, pulsar lifetime, ejecta mass, and elemental abundances; the fallback models are varied over fallback accretion rate and ejecta mass. For the abundances, we use Fe and Nd as representatives of the wind and dynamical ejecta, respectively. We simulate radiative transfer in the ejecta in either 1D spherical or 2D cylindrical spatial geometry. For the grid of 1D simulations, the mass fraction of Nd is 0, 10^{-4} , or 10^{-3} and the rest is Fe. Our models that fit the bolometric luminosity of AT 2017gfo (the kilonova associated with the first neutron star merger discovered in gravitational waves, GW170817) do not simultaneously fit the B, V, and I time evolution. However, we find that the trends of the evolution in B and V magnitudes are better matched by the fallback model relative to the pulsar model, implying the time dependence of the remnant source influences the color evolution. Further exploration of the parameter space and model deficiencies is needed before we can describe AT 2017gfo with a remnant source.

Subject headings: methods: numerical - radiative transfer - stars: neutron - supernovae: general

1. INTRODUCTION

Under the standard accretion disk paradigm for gamma-ray bursts, the outflow is powered by the release of accretion energy, typically assumed to be driven by magnetic fields. Bursts were distinguished by their duration and hardness (Kouveliotou et al. 1993). The burst duration in the accretion disk paradigm corresponds to the disk accretion timescale which, in turn, corresponds to different progenitors (Popham et al. 1999): long bursts are believed to be produced in systems where the disk can be continuously fed (e.g. collapse of massive stars) whereas short bursts are believed to be produced by compact disks (e.g. mergers of compact binaries). With these predictions for different progenitors for different bursts, theorists were able to argue for different properties of short and long bursts with respect to their locations in their host galaxies (Bloom et al. 1999; Fryer et al. 1999). Observations confirmed the distribution of locations (Fong & Berger 2013), verifying both the compact binary progenitor and the accretion disk paradigm. Neutron star/neutron star (NS/NS) and neutron star/black hole (NS/BH) mergers are the most likely compact mergers behind these short bursts.

These compact mergers have also been invoked as the source of r-process elements, the dynamically ejected material is so neutron rich that it produces a robust heavy r-process yield. Initially proposed over 4 decades ago (Lattimer & Schramm 1974), increasingly detailed studies support this as a leading source of r-process elements (e.g. Freiburghaus et al. 1999; Korobkin et al. 2012; Bauswein et al. 2013; Lippuner & Roberts 2015; Radice et al. 2016; Thielemann et al. 2017). For a review on the role of mergers in r-process production, see (Côté et al. 2018). The merger also forms a disk of high angular momentum material that drives outflows while accreting onto the central compact object. This late-time outflow ejecta is bathed in neutrinos and is likely to be less neutron rich, producing lighter (first peak r-process, iron peak) elements (Metzger &

Fernández 2014). With the broadband and multi-messenger detection of GW170817 (Abbott et al. 2017c,b,a),¹ astronomers were able to make the first definitive detection of a neutron star merger with gravitational wave (GW) measurements providing proof of (and constraints on) the merger, gamma-ray, X-ray and radio measurements of what appears to be a gamma-ray burst jet and UVOIR measurements of the merger ejecta.

With these UVOIR measurements, astronomers can, for the first time, place observed constraints on the r-process production in a neutron star merger. Moreover, these observations place constraints on nuclear composition, and hence nuclear mass models beyond the standard abundance curve (Mumpower et al. 2016). The standard models for the UVOIR emission assume a broad range of ejecta including both neutron-rich dynamical ejecta (producing heavy r-process, including lanthanides that have high optical opacities) as well as lighter elements produced by the higher electron-fraction ejecta during the accretion of the disk. The radioactive decay of these elements powers a supernova-like light curve (kilonova) and, by comparing this emission to the observations, astronomers can estimate the mass of the ejecta. The analysis of the emission to determine the exact ejecta from this merger depends upon the opacities, the opacity implementation and the distribution of this ejecta, i.e. composition, density, and velocity as a function of position (radius and angle). Models ranged from constant opacity implementations to full lanthanide opacities (utilizing a few representative isotopes) and distributions ranging from spherical mixes to multi-component ejecta models. None of the models capture all the physics and they predict a range of ejecta masses that from the observations that varies by an order of magnitude (Côté et al. 2018).

If these uncertainties were not enough, additional power sources could be augmenting the energy released from nuclear

¹ See <http://wise-obs.tau.ac.il/~arcavi/kilonovae.html> for the list of GW170817 discovery papers.

decay. As the accretion rate lowers, the radiation from the material accreting on the merged compact object is no longer trapped in the inflow and it can augment the power of the burst. In addition, there is some evidence that the merged core could be a magnetized neutron star (Piro et al. 2018). A normal, few times $10^{11} - 10^{12}$ G pulsar could also power the light curve. In this paper, we present a grid of models to study the features of these additional energy sources. The outline of our models are given in Section 2. Section 2.1 describes the 1- and 2-dimensional models used in our calculations, Section 2.2 describes our remnant sources (pulsar and accretion) and Section 2.3 describes the simulation methods used to produce kilonova light curves. With our grid of models, we produce a broad range of light curves and spectra (Section 3). Although we compare these to observations of GW170817, the intent of this project is to provide a database of spectra and light curves for upcoming observations.

2. MODEL PROPERTIES AND SIMULATIONS

2.1. Ejecta Profiles

For this project, we use both 1- and 2-dimensional outflow models. To initialize the 1D spherical ejecta, we use a slight modification to the semi-analytic, homologous solution discussed in Section 2.1.1 of Wollaeger et al. (2018),

$$v(r, t) = \frac{r}{t}, \quad (1a)$$

$$\rho(r, t) = \rho_0 \left(\frac{t}{t_0}\right)^{-3} \left(1 - \frac{r^2}{(v_{\max}t)^2}\right)^3, \quad (1b)$$

$$\mathcal{E}(r, t) = \mathcal{E}_0 \left(\frac{t}{t_0}\right)^{-4} \left(1 - \frac{r^2}{(v_{\max}t)^2}\right)^4 + E_{\text{source}}(t) \frac{\delta(r)}{4\pi r^2}, \quad (1c)$$

where r , v , ρ , \mathcal{E} , t_0 , and v_{\max} are radius, velocity, density, radiation energy density, initial time, and maximum velocity, respectively. The value of \mathcal{E}_0 accounts for the contribution of r-process nucleosynthesis and heating to the initial internal energy of the ejecta, and is the same value used by Wollaeger et al. (2018). The value E_{source} is the energy contribution from a source near the compact remnant (either from pulsar luminosity or accretion energy). We will discuss this source in more detail in section 2.2. The composition of this material is assumed to be either dominated by iron peak elements using Fe for the opacity or mostly dominated by iron peak elements with trace amounts of heavy r-process where lanthanide opacities are the most critical (with a mass fraction of 10^{-4} or 10^{-3} in Nd to represent lanthanides).

We also include a suite of models assuming a 2-component ejecta model in 2D cylindrical geometry, superimposing the spherically symmetric wind described above onto the model A dynamical ejecta of the SPH simulations by Rosswog et al. (2014) (see also Rosswog (2013)), as done by Wollaeger et al. (2018). The model A dynamical ejecta was derived from the simulation of the merger of two $1.4 M_{\odot}$ neutron stars, which produced an ejected mass of $0.013 M_{\odot}$ (Rosswog et al. 2014). For the simulations in SuperNu, this ejecta has been mapped to an axisymmetric 2D grid; hence the 3D variations around the merger axis are lost. For the wind composition, we use the same abundance options as in our 1-dimensional models, Fe. For the dynamical ejecta, we use Nd to represent a lanthanide-rich ejecta. The Model A ejecta proved to have an unobscured region permitting viewing angles where a blue

wind transient can manifest (Wollaeger et al. 2018). Consequently, the viewing-angle dependence of the light curves and spectra from this model permits the effect of the remnant luminosity to be observed at different degrees of obscurity.

2.2. Remnant Sources Enhancing the Kilonova Emission

In most kilonova light-curve models, the emission is powered by the decay of radioactive isotopes. For our models, we include this energy, but we also include the energy from an active remnant region: either pulsar or accretion luminosity. At early times, this energy is trapped in the outflowing ejecta and we can treat it as an energy source, i.e. E_{source} in Eq. (1). Let's review the features of our pulsar and accretion luminosities.

2.2.1. Pulsar Luminosity

Especially in newly formed neutron stars like the merged compact object in NS/NS binaries, pulsar emission can vary due to magnetic field restructuring, high-order field configurations, etc. Here we assume a simple dipole magnetic field source with a constant magnetic field following the pulsar luminosity formulae explored in recent studies (Lasky & Glampedakis 2016; Li et al. 2018; Piro et al. 2018). Following Lasky & Glampedakis (2016), the loss of neutron star remnant angular kinetic energy is balanced by EM dipole and GW quadrupole luminosity,

$$-I\Omega\dot{\Omega} = \frac{B_p^2 R^6 \Omega^4}{6c^3} + \frac{32GI^2 \epsilon^2 \Omega^6}{5c^2}, \quad (2)$$

where the variables are defined by Lasky & Glampedakis (2016). The pulsar luminosity is given as the electromagnetic (EM) dipole luminosity multiplied by an efficiency parameter, η , corresponding to the radiation beaming angle (Rowlinson et al. 2014; Lasky & Glampedakis 2016). Assuming GW dominated spin-down, the luminosity is (Lasky & Glampedakis 2016)

$$L_d(t) = \frac{\eta B_p^2 R^6 \Omega(t)^4}{6c^3} = L_0 \left(1 + \frac{t}{t_{\text{gw}}}\right)^{-1}, \quad (3)$$

where

$$L_0 = \frac{\eta \Omega_0^4 B_p^2 R^6}{6c^3}, \quad (4a)$$

$$t_{\text{gw}} = \frac{5c^5}{128GI\epsilon^2\Omega_0^4}. \quad (4b)$$

To explore the effect of variable remnant lifetime, we introduce a time cut-off, t_{cut} , to the luminosity,

$$L_d(t) = L_0 \left(1 + \frac{t}{t_{\text{gw}}}\right)^{-1} \Theta(t_{\text{cut}} - t), \quad (5)$$

where Θ is the unit step function.

The energy source E_{source} in Eq. (1) for our pulsar models can then be written as:

$$\dot{E}_{\text{source}} = -\frac{E_{\text{source}}}{t} + L_d, \quad (6)$$

where L_d is given from Eq. (5) and $E_d(0) = 0$. Note Eq. (6) is a simplification of the ejecta layer equations given by Metzger (2017), neglecting diffusion and changes to the inner velocity from the impulse of the pulsar luminosity. Thus we have

assumed that the pulsar luminosity is not high enough to significantly impact the morphology of the ejecta. The solution to Eq. (6) is

$$E_{\text{source}}(t) = L_0 \frac{t_{\text{gw}}}{t} \left[t' - t_{\text{gw}} \ln \left(\frac{t' + t_{\text{gw}}}{t_{\text{gw}}} \right) \right], \quad (7)$$

where $t' = \min(t, t_{\text{cut}})$; the derivative with respect to t'/t_{gw} of the term in braces is always positive, showing $E_d(t) \geq 0$.

In our simulations, $E_d(t_0)$ is added to the innermost spatial cell at a start time of t_0 . If $t_{\text{cut}} > t_0$, the pulsar is still active during the simulation, and Eq. (5) is integrated over each time step to add further energy to the innermost cell. In all simulations, we use $t_0 = 10^4$ s.

2.2.2. Fallback Luminosity

Another source of energy can come from fallback after the initial kilonova explosion. Fallback in stellar outbursts follows a simple power law with time ($\dot{m} \propto t^{-5/3}$ where \dot{m} is the fallback accretion rate and t is the time) (Chevalier 1989). The energy and mass ejected is more complex and high-resolution models have been studied in the case of supernova fallback (Fryer 2009). These models showed that roughly ~ 10 -25% of the fallback matter is re-ejected, carrying away roughly 10-25% of the accretion energy. These properties are directly applicable to the fallback in this scenario and is similar to the simple prescriptions used in the kilonova community, e.g., Li et al. (2018):

$$L_f(t) = \frac{\eta}{\eta_r} L_0 \left(\frac{\dot{m}_0}{\dot{m}_r} \right) \left(\frac{t}{t_{\text{acc}}} \right)^{-5/3} = \tilde{L}_0 \left(\frac{t}{t_{\text{acc}}} \right)^{-5/3}, \quad (8)$$

where η is again an efficiency parameter (roughly 10-25%), t_{acc} is initial time (when the luminosity begins to fall off), \dot{m}_0 is the initial accretion rate, and \dot{m}_r is a reference accretion rate. Reference values η_r and \dot{m}_r are taken to be 0.1 and $10^{-3} M_{\odot}/\text{s}$, respectively.

As with our pulsar emission, we do not modify the ejecta mass or momentum in our outflow, focusing instead on the energy injected from this accretion. This luminosity corresponds to a source energy of

$$\dot{E}_{\text{source}} = -\frac{E_{\text{source}}}{t} + L_{\text{fallback}}, \quad (9)$$

and L_{fallback} is given by Eq. (8). Assuming the luminosity is proportional to $t^{-5/3}$ back to t_{acc} , then E_{source} can be solved analytically from Eq. (9). We assume the luminosity $t^{-5/3}$ dependence holds as early as t_{acc} (which may only be true at $t \gg t_{\text{acc}}$ (Metzger 2017)). Assuming the fallback luminosity is constant before t_{acc} , the solution is

$$E_{\text{source}} = \begin{cases} \tilde{L}_0 \frac{t_{\text{acc}}^2}{2t} + \tilde{L}_0 t_{\text{acc}}^{5/3} \left(\frac{1}{t^{2/3}} - \frac{t_{\text{acc}}^{1/3}}{t} \right), & t > t_{\text{acc}} \\ \frac{t_{\text{acc}} \tilde{L}_0}{2}, & t \leq t_{\text{acc}} \end{cases}. \quad (10)$$

Unlike the pulsar models, we do not introduce a cut-off time for the source.

2.3. Methods

For the radiative transfer, we use `SuperNu` (Wollaeger et al. 2013; Wollaeger & van Rossum 2014) with tabular opacities

from the LANL suite of atomic physics codes (Fontes et al. 2015, 2017). The opacity tables used here have been described by Fontes et al (2019, submitted). For the bulk of our calculations, we use single elements to represent the material: Fe to represent the iron peak elements, Nd to represent lanthanides. For improved robustness in pre-peak luminosity kilonovae simulations with high group resolution ($N_g \gtrsim 1000$, $\lambda \in [10^3, 1.28 \times 10^5] \text{ \AA}$), we have added a more rigorous Doppler shift treatment for the diffusion optimization (Densmore et al. 2012; Abdikamalov et al. 2012; Cleveland & Gentile 2014) in `SuperNu` (Wollaeger et al 2019, in prep).

The radiative transfer is semi-relativistic and, hence, only correct to $O(v/c)$, which is a limitation that arises from the current implementation of the diffusion optimization. This issue is seemingly problematic for the pulsar kilonova models, which use low mass, high-velocity (0.3-0.45 c median) ejecta. However, a compensating phenomenon is the recession of the photosphere, which tends to relegate radiation-matter interaction (where boosting is important) to lower velocity values. For our ejecta, we may estimate the photospherical recession with

$$1 = \left(\frac{t_0^3 \rho_0 v_{\text{max}} \kappa}{t^2} \right) \int_x^1 (1 - (x')^2)^3 dx', \quad (11)$$

where κ is a grey estimate of the opacity and $x = v/v_{\text{max}}$. The integral in Eq. (11) is analytic, and the resulting expression can be solved for x given t (for instance, with Newton-Raphson iteration) or for t . For instance, for $\kappa = 0.1$ and $1 \text{ cm}^2/\text{g}$, assuming $v_{\text{max}} = 0.9c$, the time at which the photosphere reaches $v = 0.1c$ is $t \approx 0.14$ or 0.45 days, respectively. This would suggest that, on average, $O(v/c)$ radiative transfer becomes accurate on time scales relevant to observation of the wind. Additionally, for the model pulsar in particular, the spin-down emission should not be greatly impacted by the outflow speed, since the source is located at the center of the ejecta.

The radiative transfer simulations employ 64 uniform spatial cells from $v = 0$ to $v = 0.6c$, 400 logarithmic time steps from 10^4 s to 20 days, and 1024 logarithmic wavelength groups from 10^3 to $1.28 \times 10^5 \text{ \AA}$. The opacities are calculated on the same density-temperature grid as of Wollaeger et al. (2018): 17 logarithmic density points from 10^{-20} to $10^{-4} \text{ g cm}^{-3}$, 27 temperature points from 0.01 to 5 eV, and 14,900 frequency points from $h\nu/kT = 1.25 \times 10^{-3}$ to $h\nu/kT = 3 \times 10^4$ for each density and temperature. The opacity frequency grid is mapped to the radiative transfer wavelength grid by direct (unweighted) integral averaging.

Following the labeling conventions of Wollaeger et al. (2018), we call models SAFe (“semi-analytic (ejecta) with Fe”) and SAFeNd. Otherwise, we exclude the other parameter variations (mass, pulsar magnetic field, etc.) in the name, and write these out explicitly.

3. NUMERICAL RESULTS

With a range of ejecta properties (1- and 2-dimensional geometries, ejecta masses, and different compositions), we can study the role of our two energy sources, pulsars and fallback accretion. In Sections 3.1.1-3.1.3, we explore the effect of variations of remnant lifetime, elemental abundances, ejecta mass, and magnetic field strength on our pulsar kilonova models. In Sections 3.2.1-3.1.3, we vary the ejecta mass and accretion rate in our fallback kilonova models. Each variation has consequences for the observables from our models, which we discuss in the sections that follow.

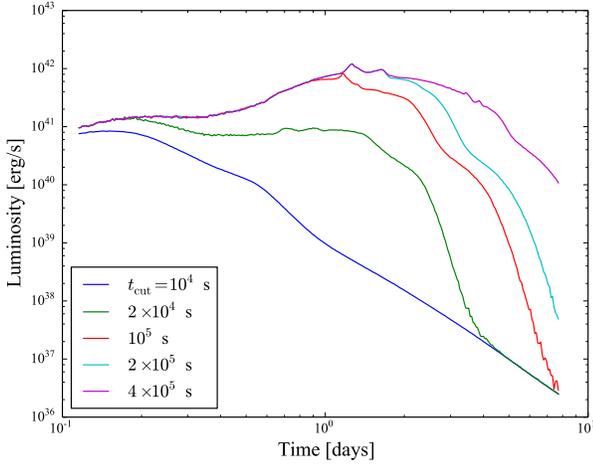


Figure 1. Bolometric luminosity versus time for the pulsar-KN models with variable remnant cutoff time. The parameters are given in Section 3.1.1. Increasing the cutoff time, t_{cut} , to values greater than 1 day does not greatly affect the peak luminosity.

3.1. Pulsar-Powered Light Curves

With our pulsar model, we have a number of free parameters. Here we study the kilonova light curves and spectra varying both the pulsar cutoff timescale and the magnetic field strength. We include both 1- and 2-dimensional geometries and vary the composition.

3.1.1. Remnant Cutoff Times

We test the effect of varying the pulsar lifetime in 1D spherical models of wind-like outflow, assuming Fe for the wind opacity. The model parameters are summarized in Table 1.

Table 1
Parameters for test of remnant lifetime.

t_{cut}	$\{0.1, 0.2, 1, 2, 4\} \times 10^5$ s
M_{ej}	$10^{-3} M_{\odot}$
$v_{\text{max}}/2$	$0.3 c$
R	1.2×10^6 cm
I	1.5×10^{45} g cm ²
Ω_0	$2\pi \times 10^3$ s ⁻¹
B	3.4×10^{12} G
ϵ	0.0035
η	1
$(\chi_{\text{Fe}}, \chi_{\text{Nd}})$	(1, 0)

The parameters for the pulsar imply $t_{\text{gw}} = 495$ s and $L_0 = 3.33 \times 10^{44}$ erg/s, similar to Li et al. (2018).

Figure 1 shows the bolometric luminosity versus time for these cutoff time variations (see model parameters above). The maximum peak luminosity is achieved for $t_{\text{cut}} \gtrsim 2 \times 10^5$ s; increasing the remnant lifetime further only affects the brightness of the tail of the light curve. For $t_{\text{cut}} \lesssim 1 \times 10^5$ s (\sim day), the peak luminosity is more sensitive to the remnant lifetime. This sensitivity can be seen in the change of the 1-day luminosity with respect to the t_{cut} .

3.1.2. Magnetic Field Strength Variations and Ejecta Mass and Composition

With a better understanding of the role of the cutoff time, we can now study the dependence of the pulsar-powered light curves on the magnetic field strength for a range of ejecta masses and two different compositions using just two cutoff times: 2×10^4 , 2×10^5 s. We vary the composition by adding a small mass fraction of Nd. Apart from the ejecta mass, magnetic field strength, and composition, the model parameters are the same as in Section 3.1.1. The model parameters are listed in Table 2.

Table 2
Parameters for test of mass and magnetic field strength variation.

t_{cut}	$\{0.2, 2\} \times 10^5$ s
M_{ej}	$\{1, 3, 10\} \times 10^{-3} M_{\odot}$
$v_{\text{max}}/2$	$0.3 c$
R	1.2×10^6 cm
I	1.5×10^{45} g cm ²
Ω_0	$2\pi \times 10^3$ s ⁻¹
B	$\{1, 10, 100\} \times 10^{12}$ G
ϵ	0.0035
η	1
$(\chi_{\text{Fe}}, \chi_{\text{Nd}})$	$\{(1, 0), (1 - 10^{-4}, 10^{-4})\}$

The parameters for the pulsar again imply $t_{\text{gw}} = 495$ s; the pulsar luminosity is $L_0 \in \{2.88, 288, 28800\} \times 10^{43}$ erg/s. Tables 3 and 4 display the luminosity at day 1 for the SAFe models with 2×10^4 and 2×10^5 s, respectively. Tables 5 and 6 show the same data for the SAFeNd models. For the lowest magnetic field strength, 10^{12} G, and lowest remnant cutoff time, increasing the mass produces an upward trend in the luminosity at day 1. For these parameters, the r-process decay energy competes with the pulsar luminosity in setting the kilonova bolometric luminosity at day 1. This is discernible in Fig. 2a, where increasing the ejecta mass can be seen to mask the pulsar peak around day 1. With a remnant lifetime of 2×10^5 s, increasing the ejecta mass from 0.001 to 0.003 M_{\odot} lowers the luminosity at day 1. The diminished luminosity at day 1 indicates delay in the emission from increased optical depth. For the higher magnetic field strengths, the luminosity at day 1 only decreases when more mass is added, resulting from the increased optical depth. The dimming and delaying of the peak luminosity for the models with $B = 10^{13}$ G and $t_{\text{cut}} = 2 \times 10^5$ s can be seen in Fig. 2c.

The introduction of the Nd mass fraction, $\chi_{\text{Nd}} = 10^{-4}$, does not appear to significantly impact the bolometric luminosity, with respect to the pure Fe models. At sufficiently late time, the bolometric luminosity is set by the r-process decay rate. However, the spectrum at later times tends to be redder; in particular for the low-pulsar luminosity, low-remnant time cutoff spectrum shown in Fig. 2c, the spectral features corresponding to the blue transient are systematically dimmer with $\chi_{\text{Nd}} = 10^{-4}$. In Fig. 2d, the brighter, longer duration pulsar luminosity appears to sustain the blue portion of the spectrum further in time.

Table 3
SAFe luminosity (erg/s) at day 1 with $t_{\text{cut}} = 2 \times 10^4$ s.

$M_{\text{ej}} (M_{\odot})$	B (G)		
	10^{12}	10^{13}	10^{14}
0.001	1.41×10^{40}	7.13×10^{41}	3.11×10^{43}
0.003	1.74×10^{40}	2.60×10^{41}	4.65×10^{42}
0.01	5.17×10^{40}	1.40×10^{41}	1.93×10^{42}

Table 4
SAFe luminosity (erg/s) at day 1 with $t_{\text{cut}} = 2 \times 10^5$ s.

$M_{\text{ej}} (M_{\odot}) \backslash B (G)$	10^{12}	10^{13}	10^{14}
0.001	9.96×10^{40}	4.55×10^{42}	8.52×10^{43}
0.003	6.48×10^{40}	2.34×10^{42}	3.67×10^{43}
0.01	8.48×10^{40}	5.36×10^{41}	1.91×10^{43}

Table 5
SAFeNd luminosity (erg/s) at day 1 with $t_{\text{cut}} = 2 \times 10^4$ s.

$M_{\text{ej}} (M_{\odot}) \backslash B (G)$	10^{12}	10^{13}	10^{14}
0.001	1.49×10^{40}	7.02×10^{41}	3.10×10^{43}
0.003	1.88×10^{40}	2.65×10^{41}	4.52×10^{42}
0.01	5.37×10^{40}	1.39×10^{41}	1.76×10^{42}

Table 6
SAFeNd luminosity (erg/s) at day 1 with $t_{\text{cut}} = 2 \times 10^5$ s.

$M_{\text{ej}} (M_{\odot}) \backslash B (G)$	10^{12}	10^{13}	10^{14}
0.001	1.00×10^{41}	4.52×10^{42}	8.38×10^{43}
0.003	6.86×10^{40}	2.31×10^{42}	3.63×10^{43}
0.01	8.57×10^{40}	5.10×10^{41}	1.81×10^{43}

We have tabulated the day 1 and day 7 bolometric luminosities and UBVRIJHK broadband magnitudes for these models in Tables 10–13 of the Appendix. The intent of tabulating the data at early and later time is to identify the effect of variations in the model on evolution of the luminosity and magnitudes. As an example, we can see that the sensitivity of the day 7 model data to remnant cutoff time depends on the pulsar magnetic strength. For instance, in Table 12, for $B = 10^{12}$ G and $M_{\text{ej}} = 0.001 M_{\odot}$, the B-band magnitude is 32.9 at day 7 for models with either $t_{\text{cut}} = 2 \times 10^4$ or 2×10^5 s. The same models with $B = 10^{13}$ G have day 7 B-band magnitudes of 32.9 and 25.3 for $t_{\text{cut}} = 2 \times 10^4$ and 2×10^5 s, respectively.

3.1.3. 2-component, 2-dimensional Pulsar Results

The morphology of the dynamical ejecta of the 2D models is “model A” used by Wollaeger et al. (2018). The wind superimposed on the model A ejecta has the properties described in Section 3.1.1 for $B = 10^{12}$ and 10^{13} G, but without a remnant cutoff time. These field strengths correspond to initial pulsar luminosities of 2.88×10^{43} and 2.88×10^{45} erg/s, respectively. Figure 3 has bolometric luminosity for the range of angular views as shaded regions for each model. The angular views are divided into 54 polar angular viewing ranges of equal solid angle. The dimmest light curves (or lowest edge of the shaded regions) correspond to views most closely aligned with the merger plane (“edge-on”), which are most obscured by Nd in the dynamical ejecta. The axial view is over an order of magnitude brighter than for the edge-on view for the model with $B = 10^{13}$ G.

Strong viewing angle dependence is also seen in broadband magnitudes. Figure 4 has plots of UBVRIJHK absolute magnitudes for each model. The viewing angle dependence is stronger for bluer bands, as expected. Before 2 days, the side view (dashed) is not sensitive to the strength of the pulsar luminosity.

3.2. Fallback Powered Light Curves

Our fallback models include a range of accretion rates, \dot{M} and ejecta masses. The variations in \dot{M} bound the typical value adopted by Li et al. (2018). These variations in M_{ej} and \dot{M} demonstrate the competition between optical depth and remnant source power, similar to the pulsar variations. Following Section 3.1, we include both 1- and 2-dimensional geometries.

3.2.1. 1-dimensional models, Accretion Rates and Ejecta Masses

For the 1D fallback models, we vary \dot{M} and M_{ej} ; otherwise the model parameters of Li et al. (2018) for the fallback luminosity source are adopted, as shown in Eq. (8). The ejecta masses, velocity, and composition are the same here as in Section 3.1.2, permitting direct comparison between these fallback models and a subset of the pulsar models. The model parameters are given in Table 7

Table 7
Parameters for test of mass and accretion rate variation.

L_0	2×10^{51} erg/s
t_{acc}	0.1 s
M_{ej}	$\{1, 3, 10\} \times 10^{-3} M_{\odot}$
$v_{\text{max}}/2$	0.3 c
η	0.1
\dot{M}	$\{1, 3, 10\} \times 10^{-3} M_{\odot}/\text{s}$
$(\chi_{\text{Fe}}, \chi_{\text{Nd}})$	$(1 - 10^{-4}, 10^{-4})$

Figure 5 displays bolometric luminosity versus time for the different accretion rates. Increasing the ejecta mass acts to dim the light curve, since the fallback source is more obscured at higher optical depth. This indicates that much of the kilonova luminosity in these models is derived from radiated accretion energy.

In the Appendix, UBVRIJHK band data at day 1 and day 7 are provided for these fallback models in Tables 14–15. The accretion rate impacts the rate of decline in the optical bands. For instance, for $M_{\text{ej}} = 0.001 M_{\odot}$, accretion rates of $\dot{M} = 0.001$, $\dot{M} = 0.003$, and $\dot{M} = 0.01 M_{\odot}/\text{s}$ drop by ~ 4.2 , 3.5, and 2.8 magnitudes in the V-band, respectively. Likewise, the R-band drops by ~ 3.9 , 3.2, and 2.7 magnitudes, respectively. The J through K-bands are relatively insensitive to the fallback source at day 1, but their brightness is more affected by day 7, where the variation with accretion rate is > 1 magnitude. In other words, at early time, the JKH magnitudes appear to be mainly set by M_{ej} , while at later time, there is a systematic increase in magnitude with increasing accretion rate.

3.2.2. 2-component, 2-dimensional Accretion Results

For the fallback models in 2D, we again use the “model A” morphology, as in Section 3.1.3. The wind superimposed on the model A ejecta has the properties described in Section 3.2.1 for $\dot{M} = 0.001$ and $0.003 M_{\odot}/\text{s}$, and with a wind mass of $0.001 M_{\odot}$. Figure 6 has bolometric luminosity for the range of angular views as shaded regions for each model. The angular views are again divided into 54 polar angular viewing ranges of equal solid angle, with edge-on views being dimmest. Relative to the 2D pulsar models, the variation with respect to viewing angle does not change as substantially when increasing the source luminosity. This is partly explained by the fallback

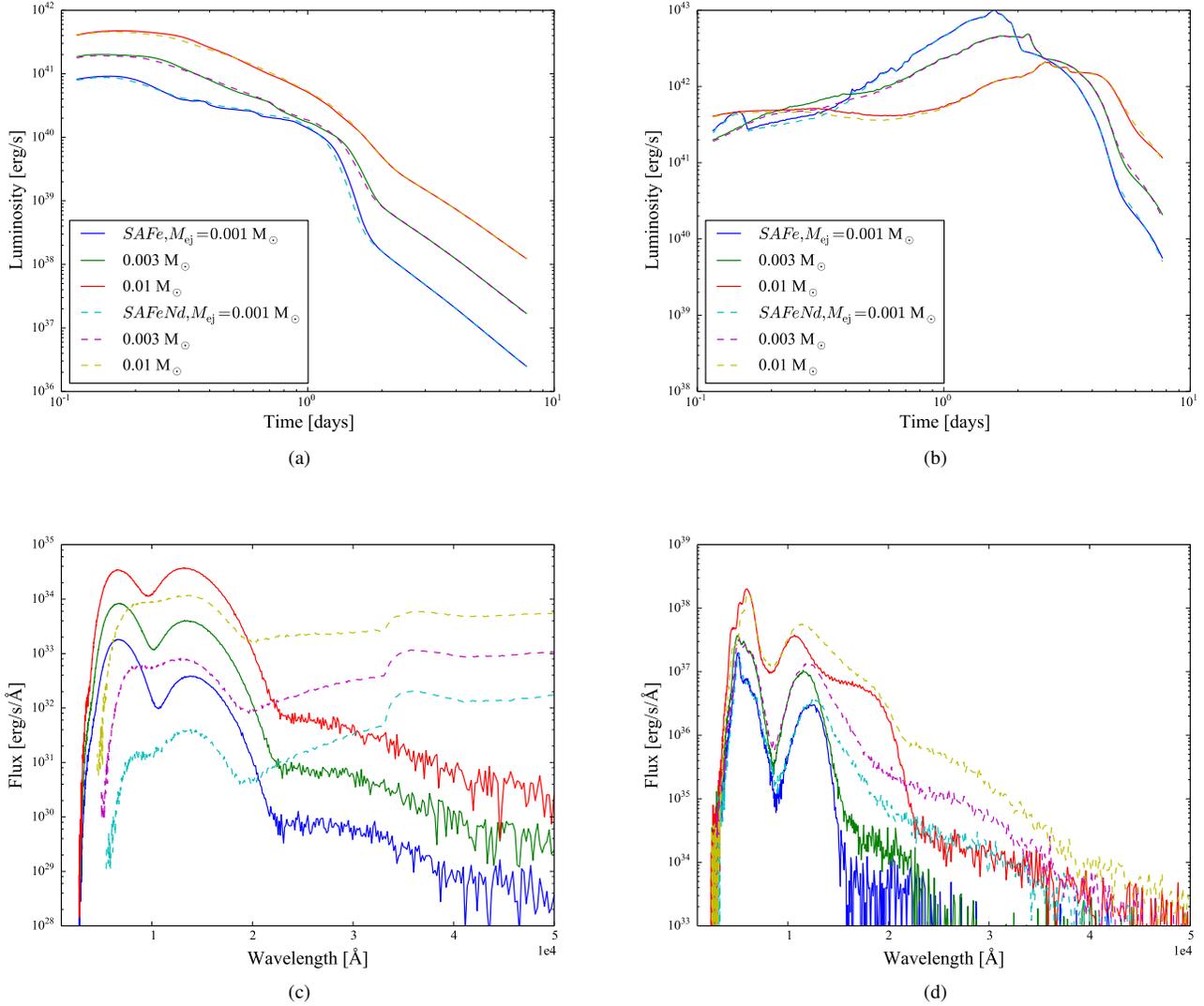


Figure 2. In Fig. 2a, bolometric luminosity versus time for each mass, for $B = 10^{12}$ G and $t_{\text{cut}} = 2 \times 10^4$ s. In Fig. 2c, corresponding spectra at day 5.45. The models with mass fraction $\chi_{\text{Nd}} = 10^{-4}$ have redder emission, despite having similar bolometric luminosity; the blue features corresponding to Fe diminish more quickly after about 2 days. In Fig. 2b, bolometric luminosity versus time for each mass, for $B = 10^{13}$ G and $t_{\text{cut}} = 2 \times 10^5$ s. In Fig. 2d, corresponding spectra at day 5.45. The persistent, brighter pulsar luminosity model seems to support the blue features of the spectrum persisting longer.

luminosity scaling linearly with accretion rate, whereas the pulsar source luminosity scales as the square of the magnetic field strength. As in the 2D pulsar models, increasing the remnant source luminosity affects the rise and decline times for near-on-axis angular views of the ejecta.

Strong viewing angle dependence is again seen in broadband magnitudes. Figure 7 has plots of UBVRIJHK absolute magnitudes for each model. The viewing angle dependence is stronger for bluer bands, as expected. Before 2 days, the side view (dashed) is not sensitive to the strength of the pulsar luminosity. Unlike the 2D pulsar models, the U-band does not significantly shift in time and peak brightness when going to a higher remnant source, which is due to the more modest increase in luminosity relative to the pulsar models.

3.3. Accretion versus Pulsar Power

With our two simple prescriptions for accretion and pulsar power sources, we produce different light curves. In Figure 8,

we compare a fallback model with $\dot{M} = 0.003 M_{\odot}/\text{s}$ to a pulsar model with $B = 3.4 \times 10^{12}$ G. It is evident that the slope of the light-curve tails are eventually set by the time-dependence of the remnant source.

3.3.1. AT 2017gfo

We can compare our two different additional power sources to observations of AT 2017gfo using the results of Li et al. (2018) and Piro et al. (2018) as a starting point. Although there is a qualitative agreement with those results, there are some differences in our conclusions. For the polar ejecta mass of $0.001 M_{\odot}$ given by Li et al. (2018), our ejecta model requires a slightly lower pulsar luminosity and somewhat higher wind outflow speed than previously published: 2×10^{44} instead of 3.4×10^{44} erg/s and $0.45c$ instead of $0.35c$. Additionally, we find that this model fits the later time bolometric luminosity better with a remnant cutoff time of about ~ 10 observer days. From the modeling perspective, some differences here

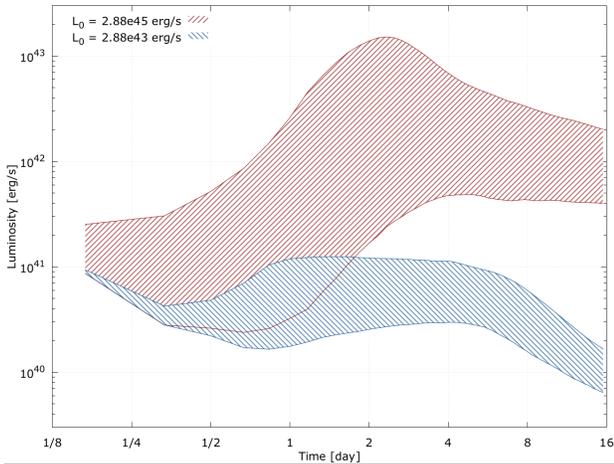


Figure 3. Bolometric luminosity versus time for the 2D pulsar kilonova models described in Section 3.1.3. The shaded regions represent the range of luminosities from the 54 angular bins that are uniform in the cosine of the polar viewing angle. Brighter luminosity corresponds to viewing bins that are more aligned with the merger axis.

include the use of multifrequency opacity tables and a detailed high-latitude r-process heating rate (corresponding to the “Wind 1” composition of Wollaeger et al. (2018)) derived from WinNet (Winteler 2012; Winteler et al. 2012). For the fallback model, similar to the polar model of Li et al. (2018), we find an accretion rate of 0.003-0.0045 M_{\odot}/s gives bolometric luminosity in the range of the observation, but with a velocity of $\sim 0.4c$ instead of $0.25c$. Parameters for the pulsar and fallback models are given in Tables and, respectively.

Table 8
Parameters for pulsar fit of AT 2017gfo.

M_{ej}	$10^{-3} M_{\odot}$
$v_{max}/2$	$0.45 c$
L_0	$2 \times 10^{44} \text{ erg/s}$
t_{gw}	495 s
t_{cut}	$2 \times 10^6 \text{ s}$
(χ_{Fe}, χ_{Nd})	$(1 - 10^{-3}, 10^{-3})$

Table 9
Parameters for fallback fit of AT 2017gfo.

M_{ej}	$10^{-3} M_{\odot}$
$v_{max}/2$	$0.4 c$
L_0	$2 \times 10^{51} \text{ erg/s}$
η	0.1
\dot{M}	$4.5 \times 10^{-3} M_{\odot}/s$
t_{acc}	0.1 s
(χ_{Fe}, χ_{Nd})	$(1 - 10^{-3}, 10^{-3})$

Figure 9 has bolometric and broadband luminosity versus time for our attempt to match AT 2017gfo with a pulsar and a fallback model. Despite the decent agreement in bolometric luminosity in Fig. 9a, it is apparent in 9b that this model produces blue emission that does not decay quickly enough with respect to the observation. Figure 9c has the light curve for the fallback model. The fallback model bolometric luminosity does not fit the data as well at intermediate times. However,

the trends in the B and V bands are much more similar to those of AT 2017gfo, relative to the pulsar model. The I-band of the fallback model is not bright enough at later times, which may be part of the disk.

4. CONCLUSIONS

We have explored several models of remnant source energy contributions to the kilonova optical/IR signal. These models are to serve as part of a larger database of models to which future observed kilonovae can be compared. Below we summarize our findings for this subset of models.

- Our attempt to fit AT 2017gfo in both bolometric luminosity and broadband light curves is unsuccessful. For a pulsar model fit of the bolometric luminosity, the broadband data is too blue in our tests, despite having an I-band in reasonable agreement with the observation. On the other hand, for the fallback model fit, the B and V-bands appear to be more consistent, but the I-band declines too rapidly. Moreover, the trend in the bolometric luminosity of the fallback model does not appear to fit well from a few days to a week.
- For the grid of models in which the remnant cutoff time, t_{cut} , was varied, we find that, for a pulsar luminosity comparable to that of Li et al. (2018), $t_{cut} \gtrsim 1$ day saturates the peak luminosity. The velocity of $\sim 0.3c$ evidently permits the effective photosphere to recede quickly enough that the pulsar energy can be uncovered on a $\gtrsim 1$ day timescale.
- As expected, increasing the pulsar luminosity and lifetime generally increases the brightness of these models. Increasing ejecta mass may either increase or decrease the kilonova luminosity. In particular, going from 0.001 to 0.01 M_{\odot} increases the r-process heating enough to mask the contribution of the low pulsar luminosity ($\sim 10^{43}$ erg/s) to the kilonova luminosity at 1 day. In contrast, with high pulsar luminosity, making the same increases to the ejecta mass lowers the luminosity at 1 day. This behavior is due to the optical depth increasing while the total r-process heating remains sub-dominant with respect to the pulsar luminosity.
- With a small contribution of Nd, the H and K bands remain brighter at day 7 for models with ejecta mass $\geq 0.003 M_{\odot}$ and magnetic field $\geq 10^{13}$ G, despite these models having similar broadband luminosity at day 1 to the versions with pure Fe.
- Following Metzger (2017); Li et al. (2018), the luminosity from fallback can significantly impact the luminosity of the kilonova model. As expected, the decay of the kilonova bolometric light-curve tail is set by the time-dependence of the remnant source. For similar source luminosity, the spectra from the fallback models are similar to those of the pulsar models; it is primarily a function of the ejecta composition. The slower decline of the luminosity tails of the pulsar models permit them to stay more luminous and have bluer spectra at later time.
- The 2D models show that the remnant source luminosity can affect viewing-angle dependence of the light curves. In particular, at early time (for instance, $\lesssim 1/2$ day for

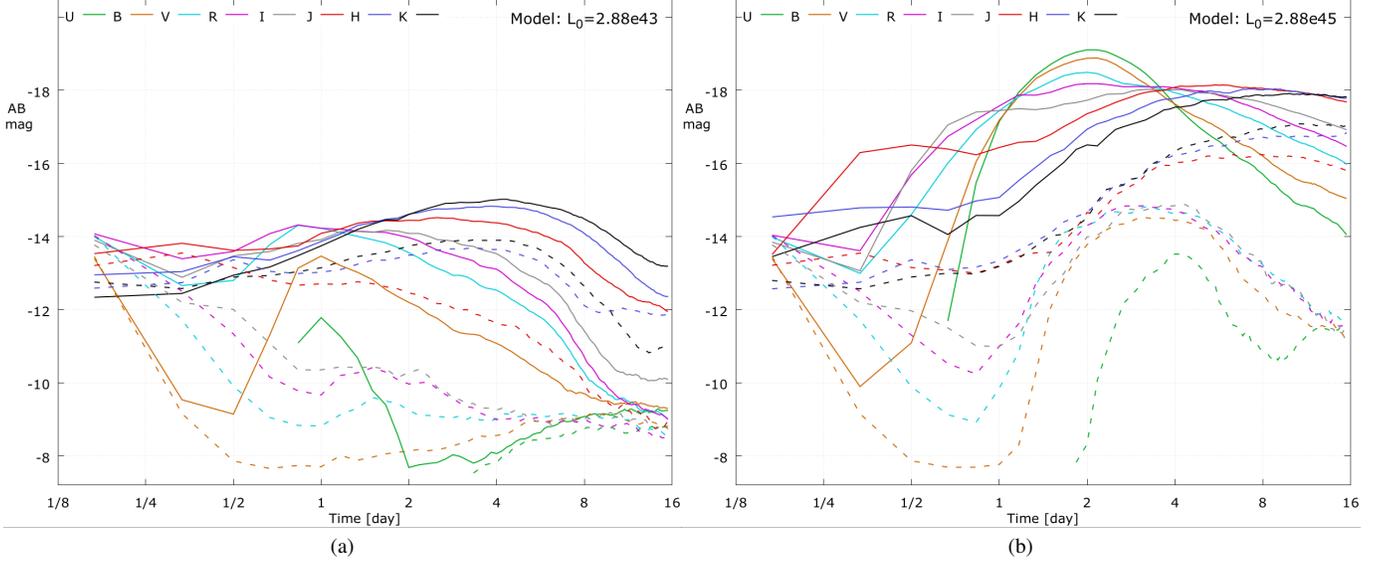


Figure 4. Top (solid) and side (dashed) angular views of the UBVRIJHK magnitudes for the 2D kilonova model described in Section 3.1.3. In Fig. 4a, the pulsar source luminosity is 2.88×10^{43} erg/s. In Fig. 4b, the pulsar source luminosity is 2.88×10^{45} erg/s.

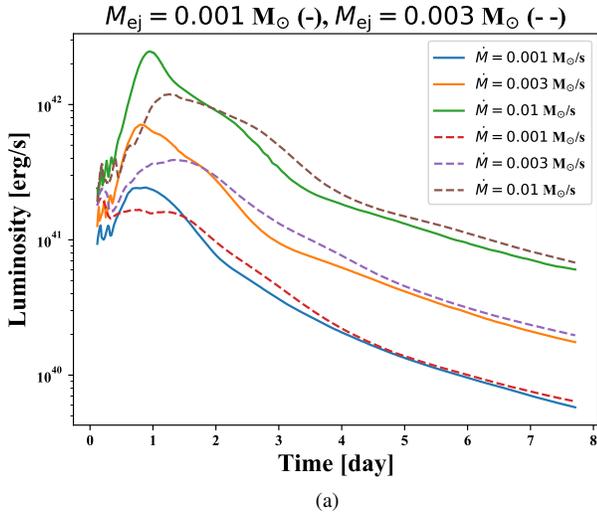


Figure 5. Bolometric luminosity versus time for the fallback kilonova model described in Section 3.2.1. The light curves were obtained under the assumption of an ejecta mass of $0.001 M_{\odot}$ (solid) or $0.003 M_{\odot}$ (dashed) and the three accretion rates.

the pulsar models explored), edge-on (off-axis) views of the ejecta produce broadband and bolometric light curves that are insensitive to the brightness of the pulsar source. As the photosphere recedes through the dynamical ejecta, the remnant source begins to contribute to the observable light curve. After 1/2 day, the time of peak magnitude for the U and B bands is shifted earlier by a few days in the higher-source energy model, while the edge on views for these bands are substantially re-brightened after the initial time.

Future work will involve attempting to refine the parameters, identifying model deficiencies, and comparing to new observations during the LIGO O3 run. The models presented in this work are a first step towards these efforts. Some concepts we have gleaned from the remnant source study, that may be

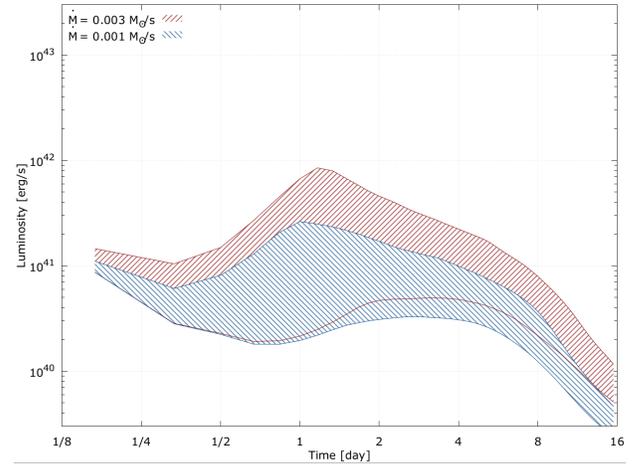


Figure 6. Bolometric luminosity versus time for the 2D fallback kilonova models described in Section 3.2.2. The shaded regions represent the range of luminosities from the 54 angular bins that are uniform in the cosine of the polar viewing angle. Brighter luminosity corresponds to viewing bins that are more aligned with the merger axis.

useful to further modeling efforts to match AT 2017gfo, are (1) the dependence of bolometric luminosity and broadband magnitudes on the time dependence of the remnant source, and (2) the competition between r-process decay heating, the remnant source luminosity, and optical depth for certain model parameters.

ACKNOWLEDGMENTS

This work was supported by the US Department of Energy through the Los Alamos National Laboratory. Los Alamos National Laboratory is operated by Triad National Security, LLC, for the National Nuclear Security Administration of U.S. Department of Energy (Contract No. 89233218CNA000001). Research presented in this article was supported by the Laboratory Directed Research and Development program of Los Alamos National Laboratory under project number 20190021DR.

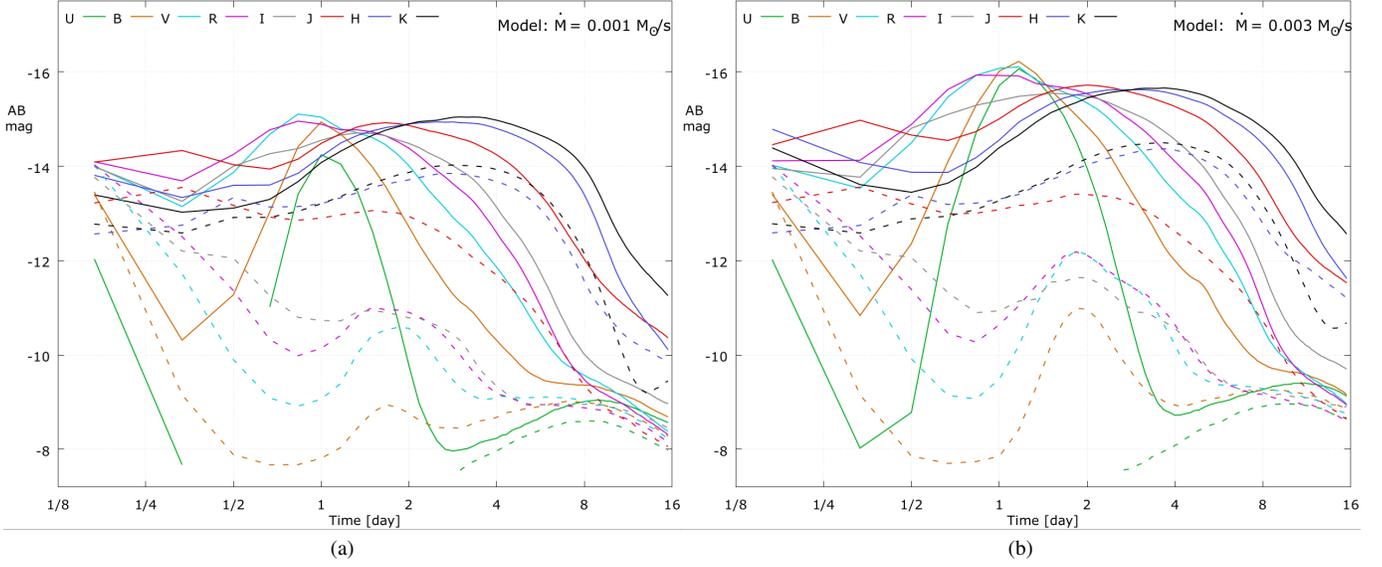


Figure 7. Top (solid) and side (dashed) angular views of the UBVRIJK magnitudes for the 2D kilonova model described in Section 3.2.2. In Fig. 7a, the fallback accretion rate is $0.001 M_{\odot}/s$. In Fig. 7b, the fallback accretion rate is $0.003 M_{\odot}/s$.

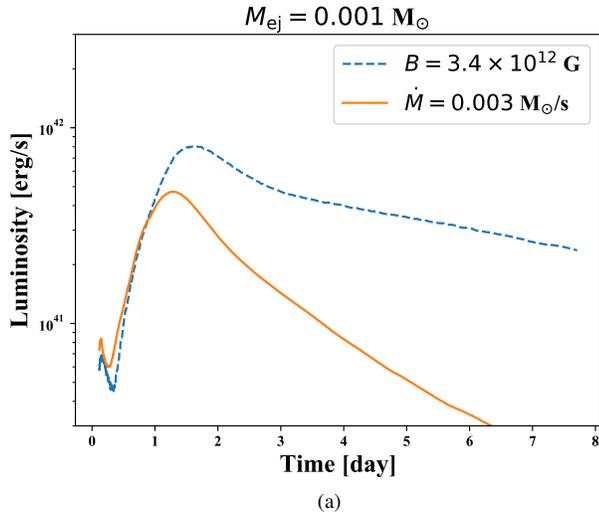


Figure 8. Bolometric luminosity versus time comparing the fallback energy source to the pulsar energy source. The figure compares a fallback model with a $\dot{M} = 0.003 M_{\odot}/s$ to a pulsar model with $B = 3.4 \times 10^{12} \text{ G}$. The slopes of the light-curve tails are eventually set by the time-dependence of the remnant source.

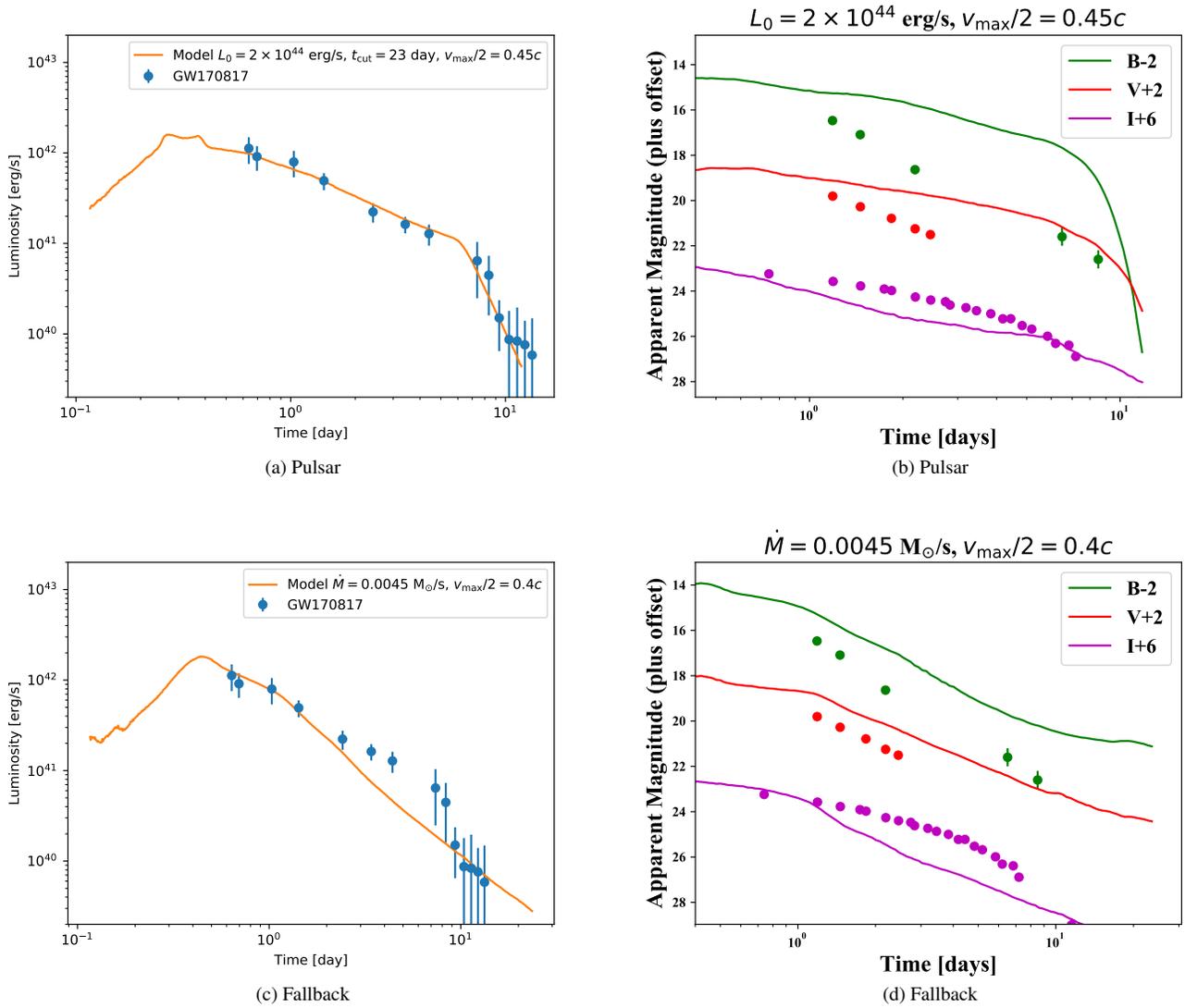


Figure 9. In Figs. 9a and 9c, bolometric luminosity versus time of the pulsar and fallback source models, respectively, described in Section 3.3.1, along with AT 2017gfo luminosity data from Piro et al. (2018). In Figs. 9b and 9d, B, V, and I broadband magnitude versus time for the pulsar and fallback source models in Section 3.3.1 along with broadband data from Troja et al. (2017). The trends in the B and V-bands of the observed data are better fit by the fallback model, which has a source luminosity that declines more rapidly (Eq. (8)).

APPENDIX

LUMINOSITY & BROADBAND MAGNITUDES FOR PULSAR KILONOVA MODELS

Table 10
Luminosity (erg/s) and 40 Mpc UBVRJHK bands at day 1 of SAFe model from Section 3.1.2.

t_{cut} (s), B (G), M_{ej} (M_{\odot})	L_{bol}	U	B	V	R	I	J	H	K
$2 \times 10^4, 10^{12}, 0.001$	1.41×10^{40}	30.1	27.0	23.3	21.6	21.6	19.7	20.0	23.4
0.003	1.74×10^{40}	29.4	26.5	23.8	21.4	21.0	19.7	19.8	21.5
0.01	5.17×10^{40}	28.1	25.3	22.9	20.5	19.8	18.9	18.5	19.3
$10^{13}, 0.001$	7.13×10^{41}	17.0	16.8	16.5	17.0	18.2	18.8	19.1	21.7
0.003	2.60×10^{41}	24.7	21.9	17.8	17.1	17.7	17.8	18.4	20.1
0.01	1.40×10^{41}	27.7	24.6	21.0	18.4	18.0	17.8	17.9	18.9
$10^{14}, 0.001$	3.11×10^{43}	13.1	14.0	12.5	11.9	13.7	15.2	17.1	17.8
0.003	4.65×10^{42}	19.5	17.2	14.8	14.0	14.7	14.6	16.0	16.7
0.01	1.93×10^{42}	22.2	19.8	17.0	16.2	16.7	14.7	15.1	15.8
$2 \times 10^5, 10^{12}, 0.001$	9.96×10^{40}	22.7	20.2	18.2	18.4	20.2	19.1	19.5	23.3
0.003	6.48×10^{40}	28.4	25.8	20.5	18.8	19.1	18.9	19.2	21.1
0.01	8.48×10^{40}	27.6	25.0	22.0	19.2	18.6	18.5	18.3	19.1
$10^{13}, 0.001$	4.55×10^{42}	15.1	15.6	14.6	14.1	15.8	17.4	18.9	19.9
0.003	2.34×10^{42}	17.4	16.6	15.2	14.9	15.5	16.4	17.7	18.8
0.01	5.36×10^{41}	24.4	21.9	17.9	16.9	17.0	16.1	16.8	17.7
$10^{14}, 0.001$	8.52×10^{43}	12.5	13.5	11.6	10.8	11.5	12.7	14.5	15.3
0.003	3.67×10^{43}	15.8	15.6	13.1	12.1	12.2	12.1	13.5	14.2
0.01	1.91×10^{43}	19.2	17.2	15.6	14.6	14.1	12.2	12.6	13.2

Table 11
Luminosity (erg/s) and 40 Mpc UBVRJHK bands at day 1 of SAFeNd model from Section 3.1.2.

t_{cut} (s), B (G), M_{ej} (M_{\odot})	L_{bol}	U	B	V	R	I	J	H	K
$2 \times 10^4, 10^{12}, 0.001$	1.49×10^{40}	30.6	27.4	23.9	21.8	21.5	19.6	20.0	22.2
0.003	1.88×10^{40}	29.6	26.7	23.9	21.4	21.0	19.6	19.7	21.0
0.01	5.37×10^{40}	28.3	25.5	22.9	20.6	19.8	18.8	18.5	19.0
$10^{13}, 0.001$	7.02×10^{41}	17.1	16.8	16.5	17.0	18.1	18.3	19.0	21.1
0.003	2.65×10^{41}	24.9	22.1	17.9	17.1	17.7	17.6	18.4	19.7
0.01	1.39×10^{41}	27.7	24.6	21.0	18.5	18.1	17.8	18.0	18.6
$10^{14}, 0.001$	3.10×10^{43}	13.2	14.0	12.6	11.9	13.7	15.2	16.9	17.9
0.003	4.52×10^{42}	19.2	17.1	14.8	14.1	14.8	14.6	16.0	16.8
0.01	1.76×10^{42}	22.5	19.8	16.9	16.2	16.8	14.6	15.2	15.8
$2 \times 10^5, 10^{12}, 0.001$	1.00×10^{41}	23.5	20.7	18.4	18.4	20.0	18.6	19.4	21.6
0.003	6.86×10^{40}	28.6	25.9	20.7	18.8	19.1	18.6	19.1	20.5
0.01	8.57×10^{40}	27.9	25.1	22.0	19.3	18.7	18.4	18.3	18.9
$10^{13}, 0.001$	4.52×10^{42}	15.1	15.6	14.6	14.1	15.8	17.3	18.7	19.7
0.003	2.31×10^{42}	17.5	16.6	15.2	14.9	15.5	16.2	17.5	18.7
0.01	5.10×10^{41}	25.2	22.5	18.1	16.9	17.1	16.1	16.9	17.6
$10^{14}, 0.001$	8.38×10^{43}	12.5	13.5	11.6	10.8	11.5	12.7	14.4	15.3
0.003	3.63×10^{43}	15.5	15.4	13.1	12.4	12.3	12.1	13.5	14.1
0.01	1.81×10^{43}	18.5	16.6	15.4	14.6	14.3	12.2	12.6	13.2

Table 12
Luminosity (erg/s) and 40 Mpc UBVRIJHK bands at day 7 of SAFe model from Section 3.1.2.

t_{cut} (s), B (G), M_{ej} (M_{\odot})	L_{bol}	U	B	V	R	I	J	H	K
$2 \times 10^4, 10^{12}, 0.001$	3.42×10^{36}	35.9	32.9	30.3	29.4	29.3	31.1	30.6	35.1
0.003	2.29×10^{37}	34.3	30.9	28.4	27.6	27.5	27.8	27.7	31.8
0.01	1.63×10^{38}	32.8	29.0	26.7	25.9	25.8	25.0	25.1	28.9
$10^{13}, 0.001$	4.82×10^{36}	35.9	32.9	30.3	29.4	29.1	30.7	29.4	30.0
0.003	2.84×10^{37}	34.3	30.9	28.4	27.6	27.5	27.6	27.3	28.2
0.01	1.78×10^{38}	32.8	29.0	26.7	25.9	25.8	25.0	24.9	27.5
$10^{14}, 0.001$	4.05×10^{40}	23.4	20.1	19.5	19.9	21.3	19.5	25.4	25.8
0.003	2.81×10^{41}	20.9	17.9	17.1	17.4	18.8	19.9	23.8	26.2
0.01	1.37×10^{42}	19.1	16.1	15.4	15.7	17.9	19.7	22.1	22.3
$2 \times 10^5, 10^{12}, 0.001$	3.44×10^{36}	35.9	32.9	30.3	29.4	29.3	31.1	30.6	34.2
0.003	2.29×10^{37}	34.3	30.9	28.4	27.6	27.5	27.8	27.7	31.6
0.01	1.63×10^{38}	32.8	29.0	26.7	25.9	25.8	25.0	25.1	28.9
$10^{13}, 0.001$	1.06×10^{40}	30.9	25.3	21.9	20.7	21.0	21.7	20.2	26.7
0.003	3.55×10^{40}	28.6	24.3	21.1	19.8	19.9	19.4	19.0	24.2
0.01	1.63×10^{41}	26.8	22.7	20.0	18.6	18.8	17.0	17.8	20.0
$10^{14}, 0.001$	4.53×10^{42}	15.5	16.4	17.7	19.7	21.9	26.5	33.5	75.7
0.003	1.01×10^{43}	14.5	15.2	17.8	18.9	20.4	25.5	26.5	31.8
0.01	2.16×10^{43}	13.3	14.3	15.3	15.9	17.3	21.9	22.8	27.0

Table 13
Luminosity (erg/s) and 40 Mpc UBVRIJHK bands at day 7 of SAFeNd model from Section 3.1.2.

t_{cut} (s), B (G), M_{ej} (M_{\odot})	L_{bol}	U	B	V	R	I	J	H	K
$2 \times 10^4, 10^{12}, 0.001$	3.40×10^{36}	-	-	44.0	40.5	38.3	34.7	34.4	33.8
0.003	2.27×10^{37}	131.9	46.8	38.3	34.6	32.7	30.4	30.5	30.6
0.01	1.63×10^{38}	86.4	38.3	33.2	30.3	28.6	26.8	27.2	27.5
$10^{13}, 0.001$	4.54×10^{36}	145.7	60.4	43.9	39.0	36.4	32.2	31.8	29.9
0.003	2.58×10^{37}	-	47.3	38.1	34.4	32.0	29.7	29.3	29.4
0.01	1.80×10^{38}	-	38.1	33.2	30.2	28.3	26.6	26.2	27.1
$10^{14}, 0.001$	4.56×10^{40}	23.5	20.2	19.3	19.8	21.1	19.2	20.7	22.2
0.003	2.93×10^{41}	21.4	18.2	17.2	17.5	18.7	18.0	19.2	20.8
0.01	1.39×10^{42}	19.7	16.4	15.5	15.7	17.7	16.5	17.8	19.1
$2 \times 10^5, 10^{12}, 0.001$	3.44×10^{36}	137.7	50.9	44.5	40.7	38.5	34.7	34.3	33.8
0.003	2.27×10^{37}	132.0	45.9	38.2	34.6	32.8	30.4	30.5	30.6
0.01	1.63×10^{38}	89.3	38.3	33.3	30.3	28.6	26.8	27.2	27.6
$10^{13}, 0.001$	1.02×10^{40}	31.3	25.7	22.2	20.8	21.1	21.8	20.1	23.7
0.003	3.53×10^{40}	29.7	24.9	21.5	20.0	19.8	19.6	18.7	22.1
0.01	1.74×10^{41}	28.7	23.8	20.8	19.0	18.5	17.1	17.2	19.5
$10^{14}, 0.001$	4.59×10^{42}	15.5	16.4	17.5	18.9	19.7	18.6	20.8	21.7
0.003	1.00×10^{43}	14.5	15.2	17.3	18.3	18.2	18.8	20.0	20.9
0.01	2.21×10^{43}	13.3	14.2	15.1	15.7	16.4	17.4	17.9	19.9

LUMINOSITY & BROADBAND MAGNITUDES FOR FALLBACK KILONOVA MODELS

Table 14

Luminosity (erg/s) and 40 Mpc UBVRIJHK bands at day 1 of SAFeNd model with fallback source, from Section 3.2.1.

$M_{ej} (M_{\odot}), \dot{M} (M_{\odot}/s)$	L_{bol}	U	B	V	R	I	J	H	K
0.001, 0.001	1.34×10^{41}	20.0	18.2	17.4	17.9	19.4	18.7	19.1	21.2
, 0.003	3.88×10^{41}	17.4	16.9	16.8	17.5	19.1	19.2	19.5	21.1
, 0.01	9.72×10^{41}	15.8	15.9	16.1	16.6	17.8	18.9	19.7	21.0
0.003, 0.001	2.56×10^{40}	24.8	22.4	18.3	17.8	19.0	17.9	18.5	20.2
, 0.003	5.75×10^{40}	21.3	19.0	16.8	17.1	18.2	17.6	18.4	20.1
, 0.01	2.53×10^{41}	18.0	16.8	16.0	16.4	17.1	17.2	18.1	19.6
0.01, 0.001	5.14×10^{40}	27.4	24.8	20.0	18.2	18.4	17.9	18.2	18.8
, 0.003	5.18×10^{40}	26.6	23.8	18.2	17.2	17.9	17.3	17.9	18.7
, 0.01	6.18×10^{40}	25.2	21.7	17.1	16.6	17.5	16.8	17.5	18.4

Table 15

Luminosity (erg/s) and 40 Mpc UBVRIJHK bands at day 7 of SAFeNd model with fallback source, from Section 3.2.1.

$M_{ej} (M_{\odot}), \dot{M} (M_{\odot}/s)$	L_{bol}	U	B	V	R	I	J	H	K
0.001, 0.001	7.66×10^{39}	27.2	23.3	21.6	21.8	22.4	20.7	22.2	23.8
, 0.003	2.44×10^{40}	25.3	21.5	20.3	20.7	21.3	19.9	20.6	23.0
, 0.01	8.61×10^{40}	22.6	19.4	18.9	19.3	20.5	19.2	20.0	22.1
0.003, 0.001	9.21×10^{39}	31.8	24.4	22.3	21.8	22.2	20.2	22.1	23.3
, 0.003	3.45×10^{40}	28.4	22.4	20.7	20.6	21.1	19.2	20.3	22.4
, 0.01	1.14×10^{41}	25.3	20.1	18.9	19.3	19.9	18.3	18.8	21.2
0.01, 0.001	1.80×10^{40}	37.8	27.4	24.5	22.4	22.4	19.9	22.4	22.8
, 0.003	5.65×10^{40}	32.3	25.3	22.8	21.1	21.1	18.7	20.4	21.8
, 0.01	1.54×10^{41}	27.6	22.3	20.2	19.6	19.6	17.6	18.2	20.7

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